



Digitization of Astronomical Photographic Plates of China and Astrometric Measurement of Single-exposure Plates

Zheng-Jun Shang (商正君)¹, Yong Yu (于涌)^{1,2}, Liang-Liang Wang (王亮亮)¹, Mei-Ting Yang (杨美婷)¹, Jing Yang (杨静)¹, Shi-Yin Shen (沈世银)^{1,3}, Min Liu (刘敏)¹, Quan-Feng Xu (徐权峰)^{1,2}, Chen-Zhou Cui (崔辰州)^{2,4,5}, Dong-Wei Fan (樊东卫)^{2,4,5}, Zheng-Hong Tang (唐正宏)^{1,2}, and Jian-Hai Zhao (赵建海)¹

¹ Shanghai Astronomical Observatory, Chinese Academy of Sciences, Shanghai 200030, China; yuy@shao.ac.cn

² University of Chinese Academy of Sciences, Beijing 100049, China

³ Key Lab for Astrophysics, Shanghai 200234, China

⁴ National Astronomical Observatories, Chinese Academy of Sciences, Beijing 100101, China

⁵ National Astronomical Data Center, Beijing 100101, China

Received 2024 January 15; accepted 2024 March 7; published 2024 May 7

Abstract

From the mid-19th century to the end of the 20th century, photographic plates served as the primary detectors for astronomical observations. Astronomical photographic observations in China began in 1901, and over a century, a total of approximately 30,000 astronomical photographic plates were captured. These historical plates play an irreplaceable role in conducting long-term, time-domain astronomical research. To preserve and explore these valuable original astronomical observational data, Shanghai Astronomical Observatory has organized the transportation of plates, taken during nighttime observations from various stations across the country, to the Sheshan Plate Archive for centralized preservation. For the first time, plate information statistics were calculated. On this basis, the plates were cleaned and digitally scanned, and finally digitized images were acquired for 29,314 plates. In this study, using Gaia DR2 as the reference star catalog, astrometric processing was carried out successfully on 15,696 single-exposure plates, including object extraction, stellar identification, and plate model computation. As a result, for long focal length telescopes, such as the 40 cm double-tube refractor telescope, the 1.56 m reflector telescope at Shanghai Astronomical Observatory, and the 1m reflecting telescope at Yunnan Astronomical Observatory, the astrometric accuracy obtained for their plates is approximately 0."1–0."3. The distribution of astrometric accuracy for medium and short focal length telescopes ranges from 0."3 to 1."0. The relevant data of this batch of plates, including digitized images and a stellar catalog of the plates, are archived and released by the National Astronomical Data Center. Users can access and download plate data based on keywords such as station, telescope, observation year, and observed celestial coordinates.

Key words: methods: data analysis – techniques: image processing – astronomical databases: miscellaneous – astrometry

1. Introduction

Before the widespread application of CCDs in the late 1980s, photographic plates played a crucial role as the primary astronomical imaging material. According to statistics, there are a total of 3 million astronomical photographic plates worldwide (Grindlay & Griffin 2012) which document celestial phenomena in observed regions that are no longer accessible. They are instrumental in conducting long-term, temporal-scale astronomical research.

In 1901, a 40 cm double-tube refractor telescope was installed in Sheshan, Shanghai, marking the beginning of astronomical photographic observations in China. Over a century, approximately 30,000 astronomical photographic plates were captured by observatories in China.

The photosensitive material of photographic plates is highly sensitive to the environment and gradually deteriorates over time. Only through digitization can astronomical plates be preserved for an extended period and conveniently utilized for scientific research. In 2000, the International Astronomical Union (IAU) established the “Plate Preservation and Digitization” working group. In 2018, the IAU General Assembly passed a resolution once again, urging global collaboration and joint efforts within the astronomical community to accelerate plate digitization. Currently, several research institutions in different countries have carried out astronomical plate digitization projects, including the DASH project at Harvard College Observatory (Simcoe et al. 2006), the APPLAUSE project jointly initiated by German observatories (Tuvikene et al. 2014), the NAROO project

Table 1
Information on Telescopes and Plates

Observatory	Telescope	Focal Length (m)	Span	Plate Size (cm)/ Field of View (degree)	Telescope Code	Plate Quantity
SHAO	40 cm double-tube refractor telescope	6.9	1901–1995	18 × 13(1.5 × 1.1), 30 × 24(2.5 × 2.0)	01	6,741
	1.56 m reflector telescope	15.6	1986–1999	20 × 20(0.7 × 0.7), 16 × 16(0.6 × 0.6), 15 × 12 (0.6 × 0.4), 30 × 24(1.1 × 0.9)	02	263
NAOC	60/90 cm Schmidt telescope	1.8	1964–1984	16 × 16(5.1 × 5.1)	03	4,265
	40 cm double-tube refractor telescope	2.0	1964–1986	30 × 30(8.5 × 8.5), 9 × 12(2.6 × 3.4), 16 × 16(4.6 × 4.6)	04	4,487
YNAO	1 m reflector telescope	13.0	1979–1993	16 × 16(0.7 × 0.7), 9 × 6(0.4 × 0.3)	05	1,017
	60 cm refractor telescope	7.5	1988	9 × 6(0.7 × 0.5)	06	9
QDO	15 cm refractor telescope(Xisha Yongxing)	0.75	1986–1987	16 × 16(12 × 12)	07	79
	32 cm refractor telescope	3.68	1964–1997	15 × 12(2.3 × 1.9), 16 × 16(2.5 × 2.5)	08	1,652
PMO	40 cm double-tube refractor telescope	3.0	1964–1997	30 × 30(5.7 × 5.7), 13 × 18(2.5 × 3.4), 12 × 9(2.3 × 1.7)	09	4,363
	60 cm refractor telescope	3.0	1954–1980	12 × 9(2.3 × 1.7)	10	1,625
	15 cm refractor telescope	1.5	1949–1976	9 × 12(3.4 × 4.6), 18 × 13(6.8 × 5.0), 10 × 15(3.8 × 5.7)	11	4,934
The total quantity of astronomical plates*						30,750

Note. *Due to missing information, the observation sources of 1315 astronomical photographic plates remain unknown.

initiated by the Paris Observatory in France (Robert et al. 2021) and the D4A project at the Royal Observatory of Belgium (De Cuyper & Winter 2005).

From 2004, Chinese astronomers called for the digitization of astronomical plates (Jin et al. 2007) and carried out preliminary preparatory work, including plate categorization, preservation and preliminary analysis of scanning accuracy (Fu & Zhao 2010; Yu et al. 2013). In 2012, China officially initiated the digitization of astronomical plates, with Shanghai Astronomical Observatory (SHAO) of the Chinese Academy of Sciences (CAS) organizing the transportation of astronomical photographic plates taken at night from various stations across the country to the Sheshan Plate Archive for coordinated preservation.

The organization of the information and the statistical work for the plates were also carried out. Between 2013 and 2017, SHAO collaborated with Nishimura Optical Manufacturing Co., Ltd. in Japan to develop a high-precision astronomical plate scanner with micrometer-level precision (Yu et al. 2017). Using this equipment, the digitization of 29,314 astronomical plates in the inventory was completed. In recent years, based on the quality of the plates, the number of exposures and the observed targets, astrometric processing of the first batch of 15,696 single-exposure plate images was carried out.

In this study, we give the first detailed introduction to the digitization of Chinese astronomical photographic plates. The focus of this study is to present a database for the China digitized astronomical photographic plates, where the astrometric measurement processing of the 15,696 single-exposure plates has been carefully performed. The paper is organized as follows. Section 2 provides an overview of the Chinese astronomical plate data. Section 3 presents the digitization status of the plates. Section 4 describes the process and results of astrometric processing for digitized images of single-exposure astronomical plates. Section 5 introduces the basic information on the database access, and finally a conclusion is presented in Section 6.

2. Overview of Chinese Astronomical Plate Data

The Sheshan Plate Archive houses 30,750 astronomical photographic plates captured by 11 telescopes from five observatories, namely SHAO, National Astronomical Observatories of CAS (NAOC), Purple Mountain Observatory (PMO), Yunnan Astronomical Observatory (YNAO) and Qingdao Observatory (QDO). These plates encompass a wide range of celestial objects, including the Moon, asteroids, comets, binary stars, variable stars, eruptive stars, radio stars, star clusters, nebulae and extragalactic objects (Jin et al. 2007). Table 1 presents the information on

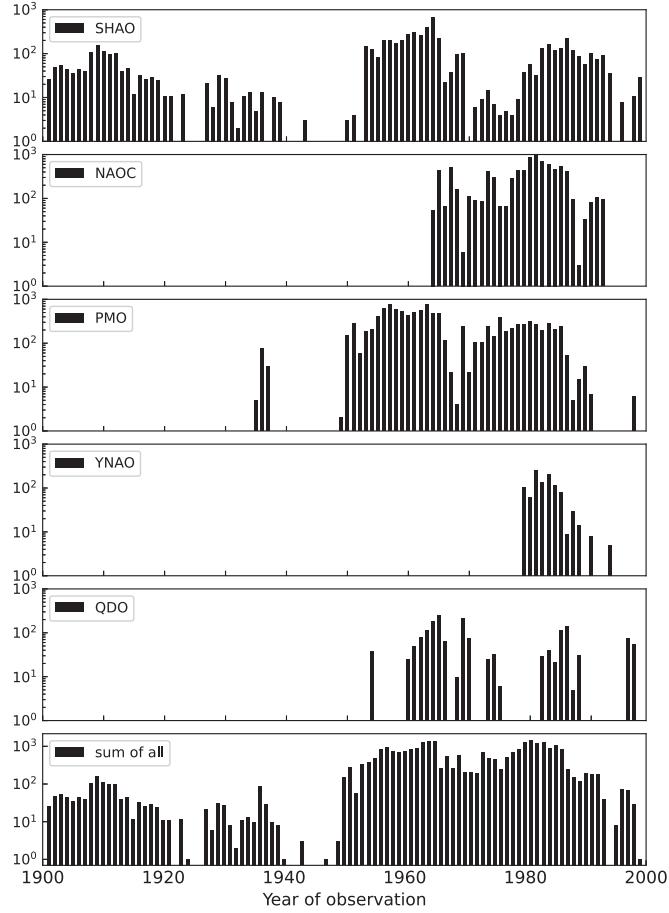


Figure 1. The distribution of the quantity of plates from each station over time. The top five panels depict the distribution of plate quantities for each station across different years, while the bottom panel represents the annual distribution of plates from all stations combined.

the telescope parameters, the span of observation years, plate sizes and the number of plates. However, due to missing record information on some plates, the sources of observation for 1,315 plates remain uncertain.

Figure 1 shows the distribution of the quantities of astronomical plates from each station over time. It can be seen that SHAO had the longest duration of plate observation, spanning 1901 to 1999. PMO had the highest number of observed plates, with a total of 10,922. Figure 2 presents a histogram of the number of plates captured by different telescopes, with each color representing the corresponding observatory. Among these telescopes, the 40 cm double tube refractor telescope (telescope No. 01), built in 1901, captured the highest number of plates, reaching 6741. Figure 3 illustrates the distribution of the pointings of all plates. It can be seen that a significant number of astronomical plates were taken along the celestial ecliptic, following the distribution of the zodiacal belt.

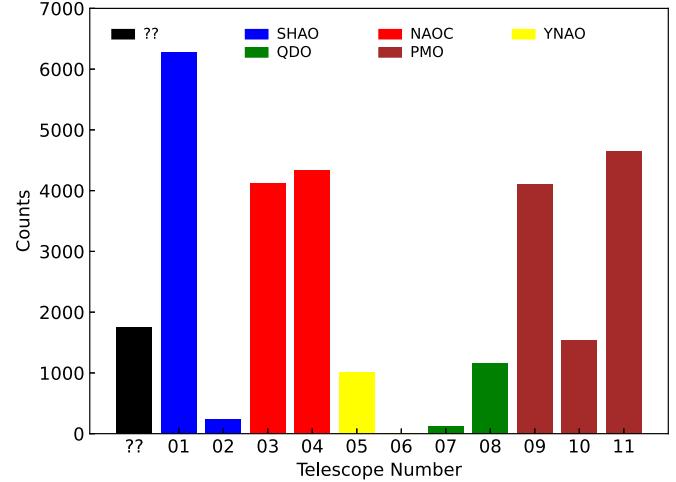


Figure 2. The distribution of the quantity of plates in the Sheshan Plate Archive by telescope. Different colors represent different observatories. The label ?? signifies plates with unidentified telescope information.

Table 2
Quantity of Astronomical Plates with Different Quality Grades

Plate Quality Grade	Quantity of Plates
1	11,475
2	9835
3	8004
4	1436

Due to their aging and various factors, there are significant differences in the quality of different plates. In 2015 June, under the guidance of experts, the plates were roughly classified according to the condition of the glass and emulsion, the extent of the mold and the quality of the plate development. Plates with intact emulsion and no mold were classified as Grade 1. Plates with minor detachment, mold or damage covering less than 25% of the surface area were classified as Grade 2. Plates with detachment, mold or damage covering less than 50% of the surface area were classified as Grade 3. Plates with severe glass damage, extensive detachment of the emulsion and significant mold were classified as Grade 4. Figure 4 provides examples of plates with different grades. Table 2 presents the number of plates in each grade. Grade 4 plates, due to severe damage, were not included in the subsequent digitization scanning process.

3. Digitization of Astronomical Plates

To avoid the inclusion of dust particles, smudges and other impurities in the scanned images, a cleaning process was performed on each plate prior to scanning. The emulsion side of the plate was cleaned using an air blower to remove surface dust. The glass side was wiped with a cotton cloth dipped in

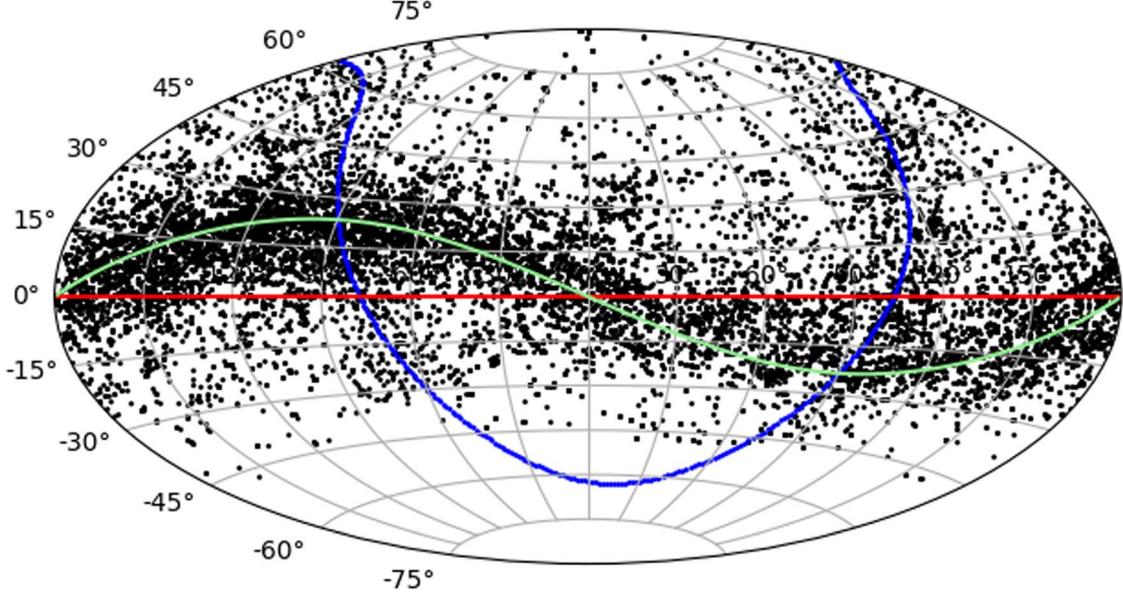


Figure 3. Distribution of plate pointing for all plates. The red line represents the celestial equator, the green line refers to the ecliptic, and the blue line signifies the galactic plane.

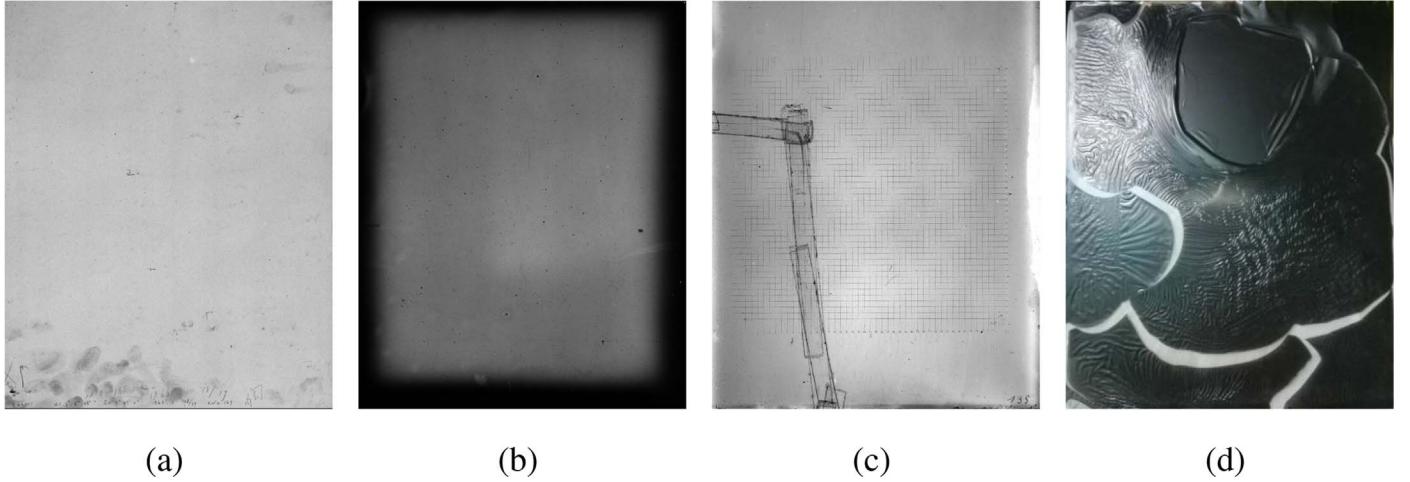


Figure 4. Examples of astronomical plates with different grades. (a)–(c): Scanned images of Grade 1–Grade 3 plates. (d): Actual photograph of a Grade 4 plate.

distilled water, followed by drying with a lint-free cloth (Yang et al. 2024).

Astronomical plate scanning was performed using a micrometer-level precision scanner (Yu et al. 2017) jointly developed by SHAO and Nishimura Optical Manufacturing Co., Ltd. in Japan, as shown in Figure 5. The scanner consists of a linear array camera, dual-sided telecentric lens, two-dimensional motion platform and an LED light source, all integrated on a granite platform. To mitigate the impact of environmental vibrations on plate scanning, the entire setup is placed on a separate base isolated from the surrounding floor.

The scanning operation room is maintained at a constant temperature and humidity ($20^\circ \text{ C} \pm 0.5^\circ \text{ C}$, $50\% \pm 5\%$) throughout the year. The scanner operates in a “line scanning” mode. Within a scanning range of $300 \text{ mm} \times 300 \text{ mm}$, the positional precision of the scans is better than $1 \mu\text{m}$, and the scanning process takes no more than 10 minutes. The scanned images are saved as 16-bit FITS files.

To facilitate the query and retrieval of scanned astronomical plate images, a standardized naming convention was applied to each scanned image after the completion of the scanning process. Taking the name “SH8201S82072001” as an

Table 3
Naming Convention for Digitized Plate Images

String	Description
SH	Two uppercase letters indicate the observatory. SH for SHAO, BJ for NAOC, ZT for PMO, YN for YNAO and QD for QDO.
82	Two digits indicate the last two digits of the year when the plate was taken. If missing, @@ is used as a placeholder
01	Two digits indicate the telescope No. If missing, && is used as a placeholder
S82072	Original plate No., if missing, ## is used as a placeholder
001	Supplementary number, defaulting to 001. It is used to label multiple plates with identical preceding information



Figure 5. Staff performing astronomical plate scanning.

example, the naming convention for image files is explained in Table 3.

4. Astrometry Processing of Single-exposure Plate Images

Based on the imaging characteristics of stellar objects on plates, the plates were classified into the following categories: single-exposure plates (18,226), multiple-exposure plates (4,632), grating observation plates (364), near-Earth moving objects (such as major planets, minor planets, comets, etc.) plates (5,778) and test observation plates (314). Figure 6 shows examples of plates from different categories, where Figure 6(a) corresponds to conventional single-exposure plates. Multiple-exposure plates refer to the repeated use of a single plate to capture multiple star images by slightly adjusting the telescope's pointing between each exposure, as demonstrated in Figure 6(b). Grating observation plates involve the use of diffraction gratings in the optical path of the telescope, resulting in multiple symmetrically distributed

Table 4
Config Values of Main Parameters in SExtractor

Parameter	Value
DETECT_THRESH	3.0
DETECT_MINAREA	5
DETECT_TYPE	CCD
BACKPHOTO_TYPE	LOCAL
MASK_TYPE	CORRECT
CLEAN	Y
FILTER	Y

diffraction orders for each stellar object, as verified in Figure 6(c). Near-Earth moving object plates are obtained by tracking moving targets with the telescope, where the target objects appear as circular shapes, while stellar objects exhibit elongation, as displayed in Figure 6(d). Test observation plates are captured to adjust telescope pointing, focal length, exposure time and other parameters, as depicted in Figure 6(e).

In this paper, we present the astrometric processing on single-exposure plates. Processing of other complex plates will be carried out in follow-up work.

Astrometric processing of plates involves several steps, including star extraction, reference star matching, plate parameter model computation and celestial coordinate calculation. The workflow is illustrated in Figure 7. First, SExtractor (Bertin & Arnouts 1996) is used to obtain the measured coordinates and instrumental magnitudes of all candidate stellar objects on the plates. The parameter settings for the SExtractor software, tailored to the specific plates, are listed in Table 4. In the output parameters, we use XWIN_IMAGE, YWIN_IMAGE as measured coordinates of the objects, and MAG_AUTO as the instrumental magnitudes of the objects.

Then, the preliminary calibration of the detected objects is performed using Astrometry.net (Lang et al. 2010). If the calibration is successful, the reference stars obtained from Astrometry.net are used to calculate the parameters of the plate model and determine the celestial coordinates of all stars. By matching with the Gaia Data Release 2 (DR2) catalog (Prusti et al. 2016; Brown et al. 2018), additional reference stars are identified. This process is repeated until a

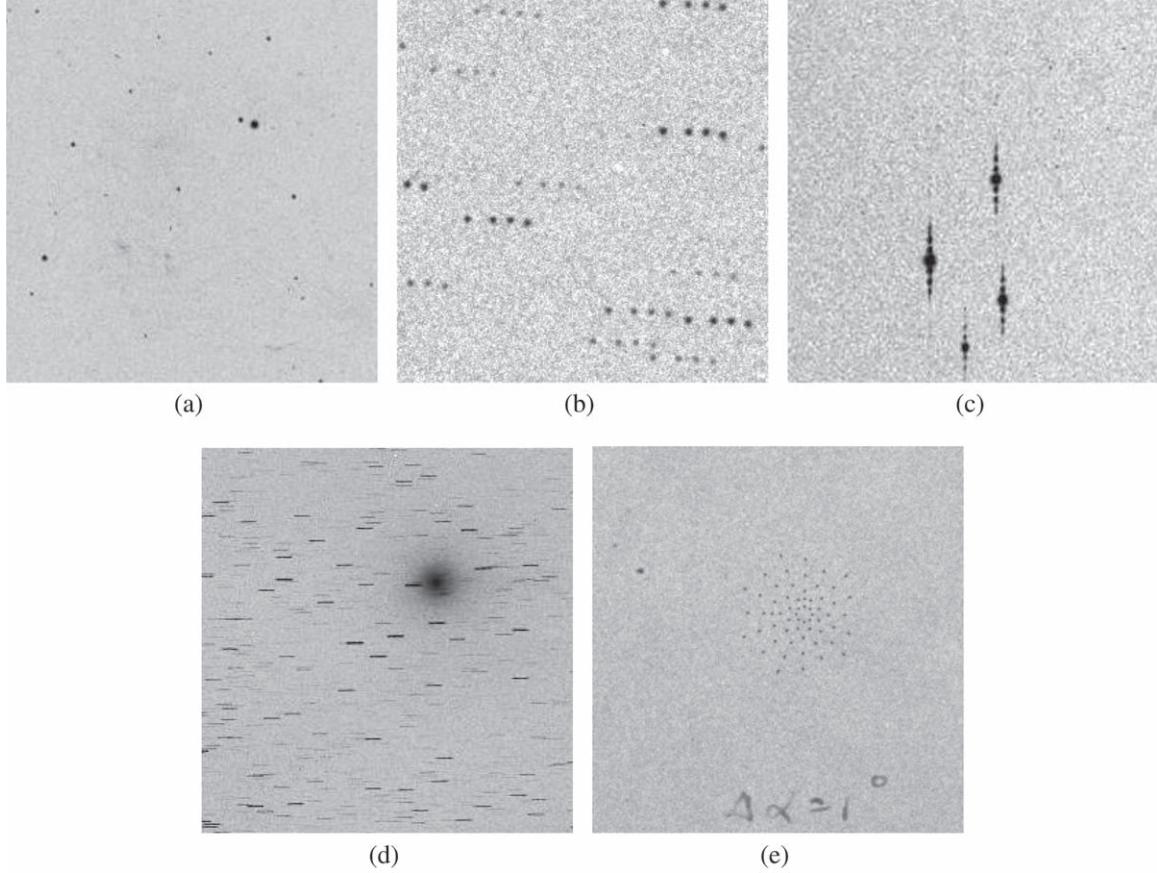


Figure 6. Examples of plates of different categories. (a) represents a single-exposure plate, (b) represents a multiple-exposure plate, (c) represents a grating observation plate, (d) represents a plate capturing near-Earth moving objects and (e) represents a test observation plate.

Table 5
Proportion of Stellar Objects in the Targets Before and After SVM Labeling (Using 10 Plate Images as an Example)

Plate Name	n_1	n_2	n_2/n_1	n_3	n_4	n_4/n_3
SH6201V62027001	33,957	4,400	0.13	5,607	2,931	0.52
SH6201V62036001	57,650	6,302	0.11	11,902	4,979	0.42
ZT7509III267001	57,30	716	0.12	643	336	0.52
SH6201V62021001	57,375	5,039	0.09	6,679	3,718	0.56
ZT6111NO65001	2,294	176	0.08	234	100	0.43
SH5901G59041001	9,986	1,191	0.12	2,526	1,059	0.42
SH5701CL57039001	21,520	1,948	0.09	3,907	1,617	0.41
SH0501145001	62,407	8,181	0.13	14,779	6,668	0.45
SH8201CL82006001	88,887	4,971	0.06	5,915	3,681	0.62
BJ8303SD1867001	238,487	11,166	0.05	10,067	9,487	0.94

Note. n_1 is the number of targets extracted by SExtractor. n_2 is the number of targets matched to Gaia DR2. n_3 is the number of targets labeled as stellar objects by SVM. n_4 is the number of true stellar objects (matched to Gaia DR2) among the targets labeled as stellar objects by SVM.

stable match result with Gaia DR2 reference stars is achieved. For some plate images with many scratches, mold spots and other artifacts, direct calibration using Astrometry.net may fail. In such cases, we employ the Support Vector Machine

(SVM) method to filter out non-stellar objects from the extracted objects, which can improve the success rate of complex plate calibration for these challenging images. For more details, see Section 4.1.

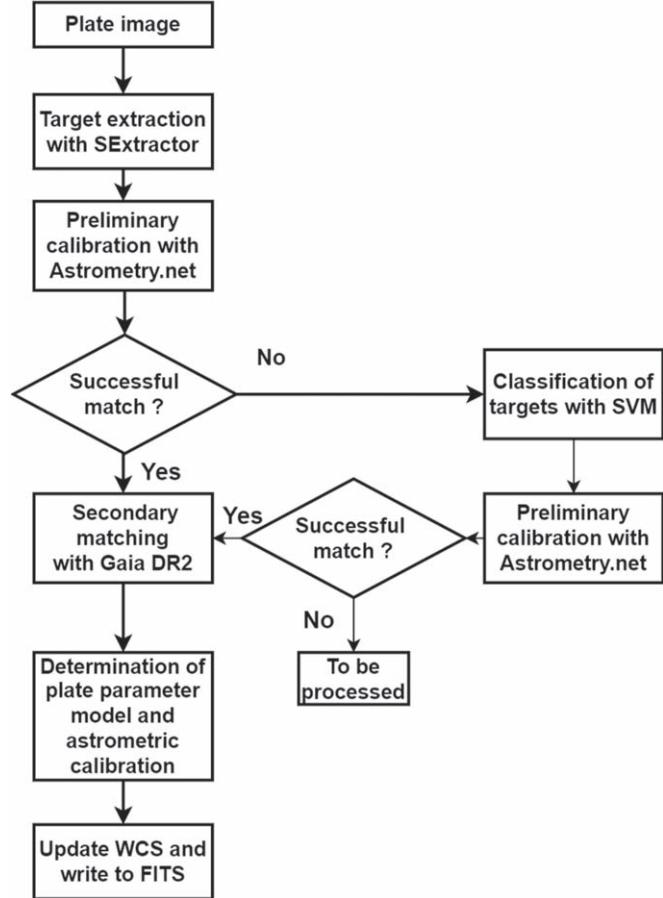


Figure 7. Flow chart of astrometric calibration for single-exposure astronomical plate images.

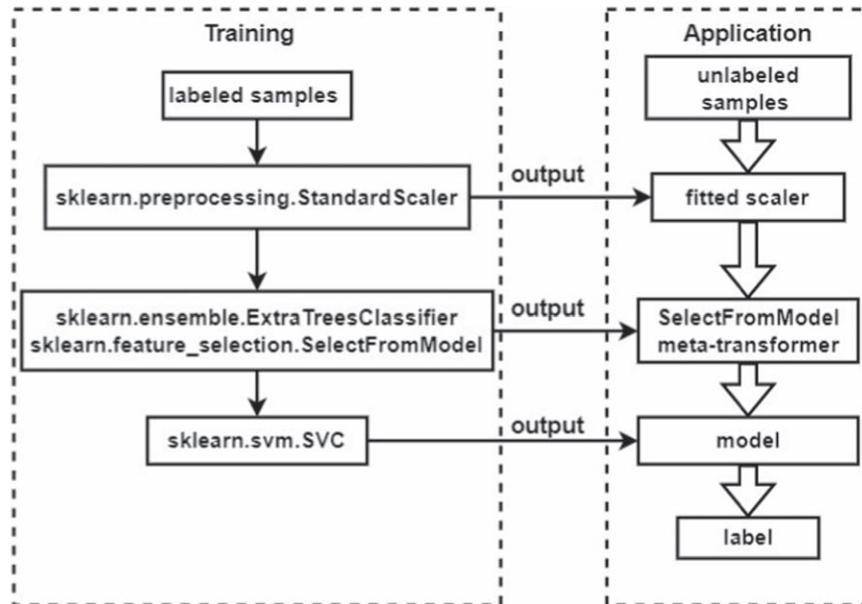


Figure 8. Training and application workflow of the SVM model.

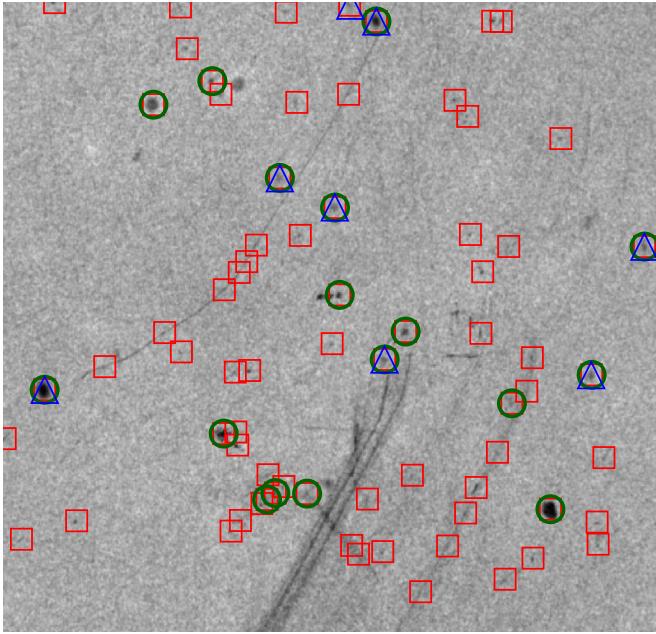


Figure 9. Illustration of target classification for plate SH6201V62036001. The red squares represent the targets extracted by SExtractor, the green circles correspond to the targets classified as stellar objects by SVM and the blue triangles signify the targets matched with Gaia DR2.

The large physical size of astronomical plates (greater than $6\text{ cm} \times 9\text{ cm}$) makes their imaging susceptible to various factors such as optical distortions of the field of view, non-vertical alignment between the plate and telescope optical axis, poor atmospheric refraction, etc. Additionally, nonlinearity in response of the plate to light intensity and asymmetric imaging optical paths can lead to different center offsets for images of stars with different magnitudes, also known as the magnitude equation (Girard et al. 1998). These combined factors contribute to the imaging patterns of the plate. Therefore, through sampling and analysis of observational data from different telescopes, we have determined the appropriate calibration models for the respective telescope plate images. More details and analysis results can be found in Section 4.2. After combining the above methods, 15,696 single-exposure plates were astrometrically calibrated, and the calibration results are presented in Section 4.3.

To facilitate the query and retrieval of astronomical plate images, for each plate image, after completing the model computation for each plate, the model parameters are converted into World Coordinate System (WCS) parameters in the TAN-SIP format (Shupe et al. 2005). These parameters include the celestial coordinates of the plate center, linear terms and nonlinear terms, which are then written into the header of the image FITS file.

4.1. Enhanced Stellar Target Selection with SVM

Due to the interference of scratches, mold spots and impurities, a total of 4714 single-exposure plate images cannot be used directly for initial reference star calibration using Astrometry.net. For these plates, we employ SVM for the classification of “stellar objects” and “non-stellar objects”. SVM is a machine learning method that aims to find an optimal hyperplane in an n -dimensional feature space by minimizing the objective function and effectively distinguishing between targets. SVM offers advantages such as simplicity and strong generalization capabilities (Noble 2006; Chauhan et al. 2019), which have been widely used in the classification of celestial objects (Zheng et al. 2005; Bu et al. 2014; Ma et al. 2016).

An SVM model was established using data from plates with completed astrometric calibration. SVM is implemented by calling scikit-learn (Pedregosa et al. 2011), and the basic process is diagrammed in Figure 8. In the first step, the parameters from the SExtractor output, excluding the constant and positional parameters, were selected as sample features, including the parameters related to MAG, AREA, ERROR and shape. Each extracted object consists of a data sample in the SVM algorithm. In the second step, the samples are standardized to ensure that each feature has a similar impact on the model’s classification ability, avoiding individual features dominating the model’s classification ability. In the third step, feature selection is performed by constructing an extremely randomized forest classifier to rank the features based on their importance. Features that exceed a threshold are selected for subsequent training and application. The selected feature parameters and sample labels are then used for model training. Finally, the model is applied to plates where the initial calibration using Astrometry.net has failed, and the extracted objects are labeled as “stellar objects” or “non-stellar objects”.

Randomly taking 10 plates as an example, Table 5 shows the proportion of true stellar objects before and after SVM classification. n_2/n_1 represents the proportion of true stellar objects among objects directly extracted from SExtractor. n_4/n_3 represents the proportion of true stellar objects among targets classified as “stellar objects” by SVM. It can be observed that the proportion of true stellar objects significantly increases after the SVM classification. Using the plate SH6201V62036001 as an example, Figure 9 displays the SExtractor extraction results, the SVM classification results and the results from matching with the Gaia DR2 star catalog.

After SVM target selection, Astrometry.net was used again for initial calibration, and 2,184 calibrated results were successfully obtained from 4,714 plate images. Therefore, we finally had 15,696 single-exposure plates (18226 - 4714 + 2184) that passed this initial step of astrometric processing,

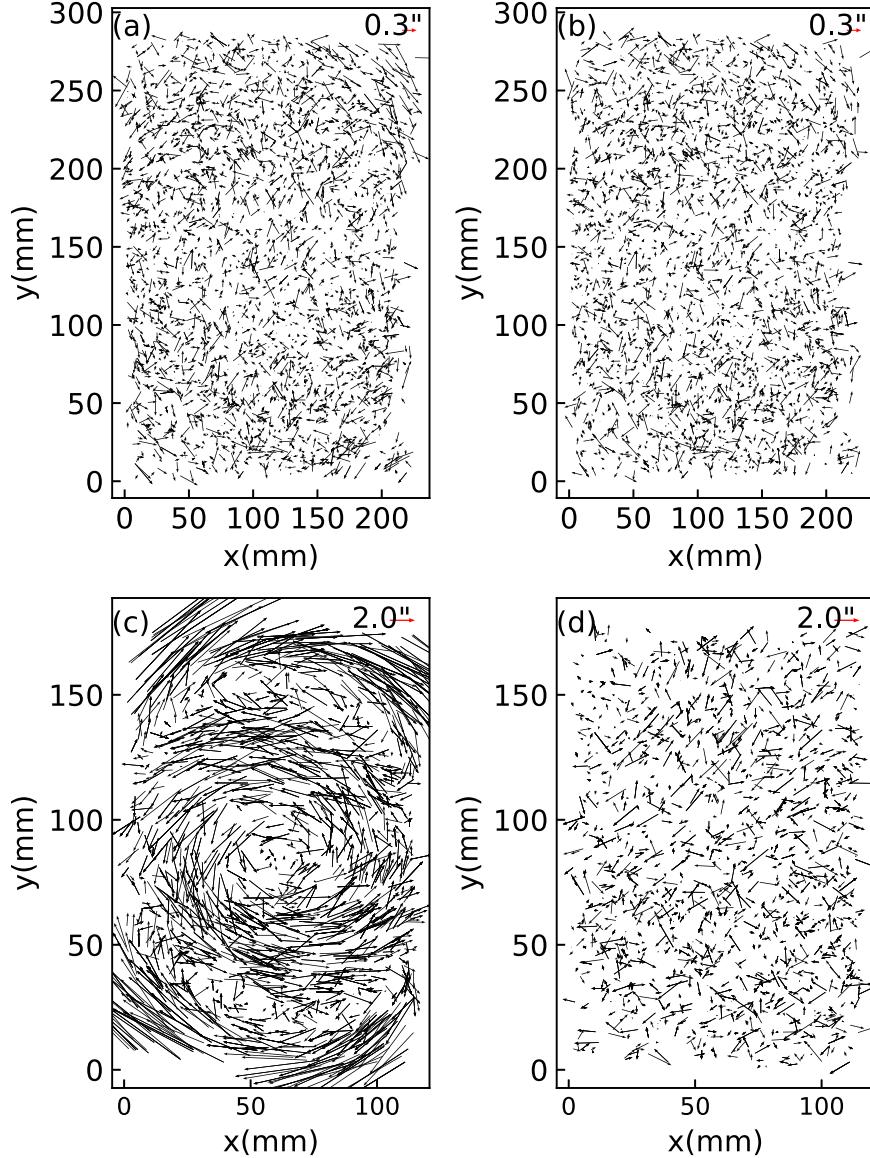


Figure 10. Distribution of reference star residuals on the image. (a) and (b) represent the case of plate SH6301V63096001 taken with telescope No. 01 at SHAO, using 1st- and 2nd-order model reduction, respectively. (c) and (d) represent the case of plate ZT5011T123001 taken with telescope No. 11 at PMO, using 2nd and 3rd-order model reduction respectively.

which constitutes the astronomical photographic plate database presented in this study.

4.2. Selection of Plate Parameter Model

The plate parameter model is used to describe the relation between the ideal coordinates and measured coordinates of stellar objects on the plates, including translation, rotation, scale and distortion. To determine the appropriate form of the plate parameter model, a random sample of 100 plates from each telescope was taken. The plates were individually reduced using conventional 1st, 2nd, 3rd and 4th-order

models (i.e., 6, 12, 20 and 30-parameter models, respectively), and the standard deviation of the residuals concerning the Gaia DR2 reference stars was examined for each plate model reduction. The best model is chosen based on the criteria of having the lowest order, lowest standard deviation and randomly distributed residuals. Table 6 presents the standard deviation of the reference star residuals for each telescope using different models. It can be seen that for the plates taken with telescope No. 01 (40 cm double refractor telescope at SHAO), the standard deviation of the 2nd-order model is comparable to that of higher-order

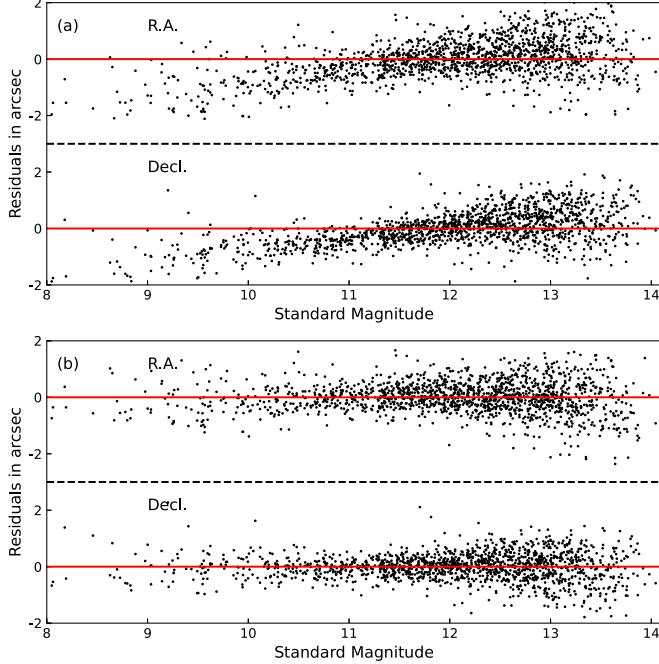


Figure 11. Distribution of residuals of reference stars with magnitude. Data are from plate ZT5011T123001 taken with telescope No. 11 at PMO, with the reduction of a 3rd-order model. The upper panels are residuals of R.A. and the lower ones are of Decl. (a) shows the distribution of residuals without magnitude equation correction. (b) displays the distribution of residuals with magnitude equation correction.

models. For plates taken with other telescopes, the standard deviation of the 3rd-order model is comparable to that of higher-order models. Taking plate SH6301V63096001 from telescope No. 01 and plate ZT5011T123001 from telescope No. 11 as examples, Figure 10 shows the distribution of the reference star residuals in the image. It can be seen that the residual distribution is predominantly random. Therefore, except for the plates taken with telescope No. 01 at SHAO, which are reduced using a 2nd-order plate parameter model, plates from other telescopes are reduced using the 3rd-order plate parameter model.

A direct consequence of the magnitude equation is that the position offsets are related to stellar magnitudes. Taking the plate ZT5011T123001 as an example, Figure 11(a) illustrates the distribution of the residuals of reference stars as a function of their magnitudes. Note that the magnitude used here is the g magnitude of the reference stars in the Gaia DR2 catalog. It can be seen that the distribution exhibits significant systematic patterns, with a maximum value of up to $2''$. To mitigate the impact of this factor, we introduced a magnitude correction term in the reduction of the plate parameter model, which is $p_1 \cdot \text{mag} + p_2 \cdot \text{mag}^2$, where mag represents the magnitude of the reference star, and p_1 and p_2 are the coefficients to be determined. Figure 11(b) displays the distribution of the

Table 6
Standard Deviations of Reference Star Residuals Under Different Order Plate Parameter Models

Telescope No.	Standard Deviation ($''$)				Mean Quantity of Reference Stars
	1st-order model	2nd-order model	3rd-order model	4th-order model	
01	0.138	0.125	0.123	0.122	1,996
02	0.149	0.132	0.097	0.097	768
03	0.525	0.439	0.349	0.345	27,818
04	3.386	2.098	0.593	0.592	65,228
05	0.220	0.213	0.107	0.105	454
08	0.383	0.334	0.308	0.306	2,184
09	0.718	0.600	0.271	0.271	21,116
10	0.761	0.721	0.593	0.587	1,559
11	2.790	2.265	0.599	0.596	3,169

residuals of the reference stars after incorporating the magnitude correction term. The correlation between the residuals and stellar magnitudes is significantly reduced.

4.3. Astrometric Calibration Results

Using Gaia DR2 as the reference star catalog, astrometric calibration was conducted on a total of 15,696 single-exposure plates, obtaining the celestial coordinates of the stellar objects on each plate in the International Celestial Reference System (ICRS). Figure 12 presents the distribution of standard deviation of the residuals of the reference stars on each plate by telescope. The standard deviation reflects the astrometric accuracy of the stars on the plates, mainly influenced by the focal length of the telescope and possibly by the condition of plate preservation. For telescopes with longer focal lengths, such as the 1.56 m reflecting telescope (No. 01) and the 40 cm double tube refracting telescope (No. 02) at SHAO, and the 1m reflecting telescope (No. 05) at YNAO, the accuracy distribution of the plates ranges from $0.^{\circ}1$ to $0.^{\circ}3$. For telescopes such as the 60/90 cm Schmidt telescope (No. 03) at NAOC, the 32 cm refracting telescope (No. 08) at QDO and the 40 cm double-tube refracting telescope (No. 09) at PMO, the distribution of the accuracy of the plates ranges from $0.^{\circ}3$ to $0.^{\circ}5$. Finally, for the 40 cm double tube refracting telescope (No. 04) at NAOC, and the 60 cm reflecting telescope (No. 10) and the 15 cm reflecting telescope (No. 11) at PMO, the distribution of the accuracy of the plates ranges from $0.^{\circ}4$ to $1.^{\circ}0$.

Figure 13 shows the distribution of the standard deviation of positional residuals of reference stars by telescope with magnitude. The magnitudes used here are also the g magnitudes of the reference stars in the Gaia DR2 catalog. It can be seen from the figure that the distribution is different for different telescopes. However, there is a general trend that the position residuals of the reference stars first

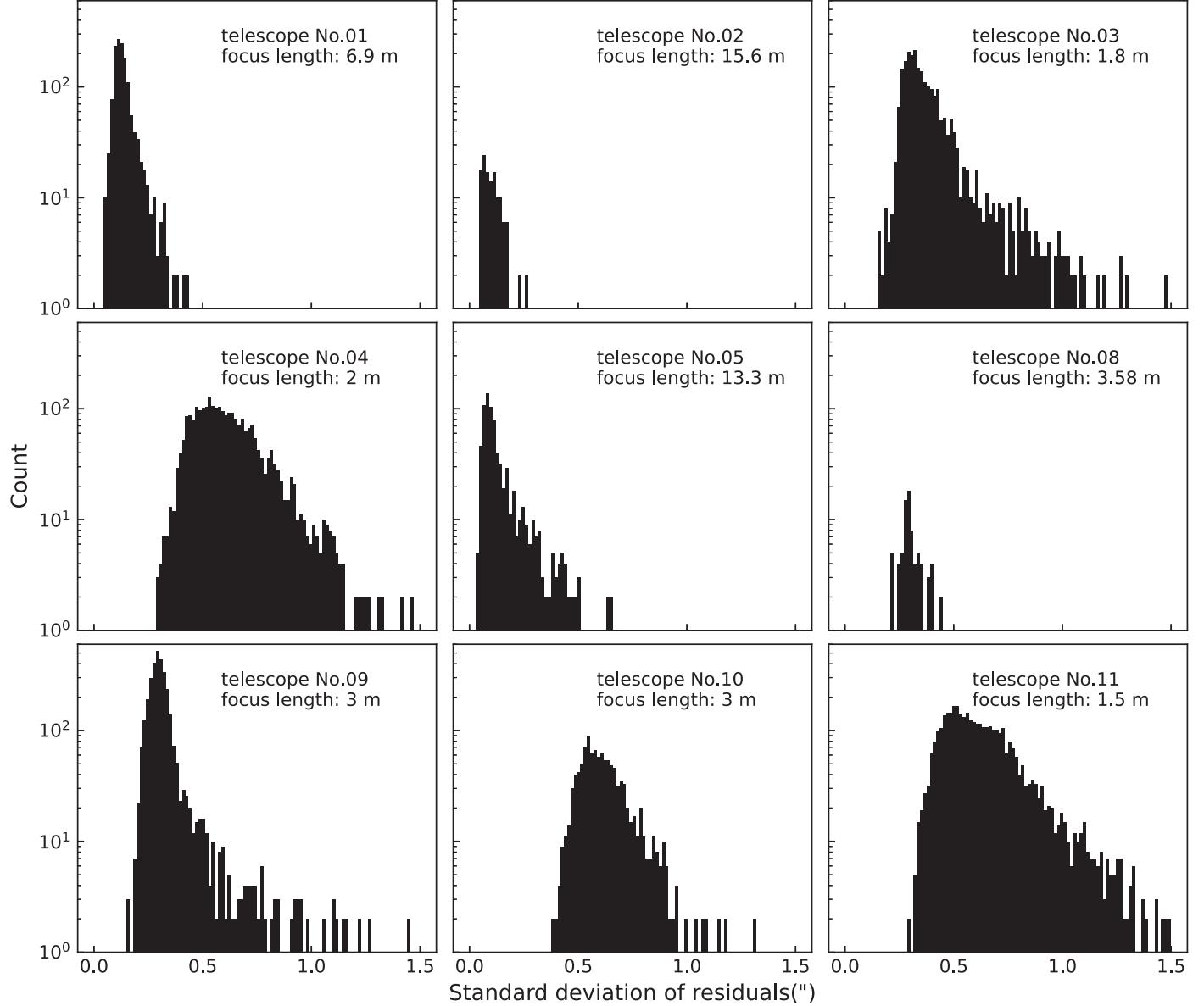


Figure 12. Distribution of standard deviation of the positional residuals of reference stars on each plate by telescope. None of the plates from telescope Nos. 06 and 07 were astrometrically calibrated successfully, so there are no data from telescope Nos. 06 and 07.

decrease with magnitude and then increase with magnitude, and objects with medium magnitude have the highest position accuracy. The large positional residuals of bright stars may come from inaccurate centering due to saturation, while the large positional residuals of weak stars may arise from inaccurate centering due to excessive noise. Medium-brightness stars are less affected by saturation and noise, so they are most accurately centered, and thus have less positional residuals than bright and weak stars. Figure 14 displays the intensity distributions of stars with different brightnesses.

Photometric calibration of a photographic glass plate is a complex task in itself, and this is compounded by the missing filter information for the plates. Here we have only roughly fitted the instrumental magnitudes to the g magnitudes of the reference stars in the Gaia DR2 catalog with the *rlowess* algorithm (Cleveland 1979). Detailed photometric calibration is still in progress and will be presented in a subsequent study.

5. Database

In this study, we release the digitized images of all 15,696 single-exposure plates and their corresponding stellar catalogs.

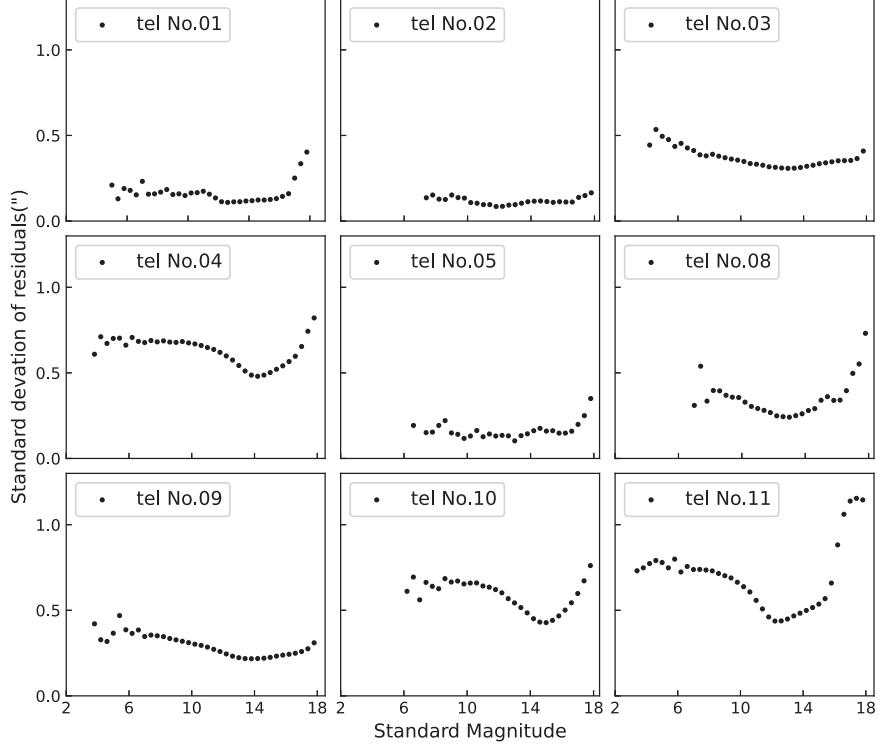


Figure 13. Distribution of the standard deviation of the positional residuals of the reference stars by telescope with magnitude.

The data release of this batch of plates, named the China Astronomical Plates Data Release, are archived at the National Astronomical Data Center (NADC) with doi [10.12149/100742](https://doi.org/10.12149/100742) and CSTR 11379.11.100742. Users can access the data through the website <https://nadc.china-vo.org/res/r100742/>, where the interactive search and download functions on the plates and stellar catalogs are provided.

6. Conclusion

To preserve and explore valuable astronomical information on photographic plates, SHAO has organized the transportation of astronomical photographic plates taken at night from various observatories nationwide to the Sheshan Plate Archive for unified preservation. In addition, efforts have been made to compile and analyze plate information. Currently, the Sheshan Plate Archive houses astronomical plates from 11 telescopes, including those from SHAO, NAOC, PMO, YNAO and QDO, totaling 30,750 plates. These plates span nearly 100 yr of astronomical observations.

Based on factors such as degree of damage, mold affected areas and the plate development process, a quality grading system was implemented for the plates. Building upon this grading, cleaning and digital scanning of the plates were carried out, resulting in a final collection of 29,314 digital plate images. Astrometric calibration has been performed

on 15,696 single-exposure plates, which involved processes such as object extraction, star matching with the Gaia DR2 catalog and plate parameter model computation. The results indicate that for long-focus telescopes, such as the 40 cm double-tube refracting telescope and the 1.56 m reflecting telescope at SHAO, and the 1 m reflecting telescope at YNAO, the astrometric accuracy of the plates ranges from approximately $0.^{\circ}1$ – $0.^{\circ}3$. For other medium- and short-focus telescopes, the astrometric accuracy of the plates ranges from $0.^{\circ}3$ to $1.^{\circ}0$.

In the future, we will try to solve the astrometric calibration of the remaining 2,580 single-exposure plates, and find the optimal method of photometric calibration for all plates. For the multi-exposure plates, plates capturing near-Earth moving objects and grating observation plates, we will proceed to carry out their astrometric and photometric calibration.

Data relating to this batch of 15,696 single-exposure plates, including digitized images of the plate with observation information and WCS parameters, as well as the scanned images of plate envelopes, are available at the website of NADC (<https://nadc.china-vo.org/res/r100742/>). Users can query, retrieve and download plate data based on keywords such as observatory, telescope, observation year and observation sky coordinates.

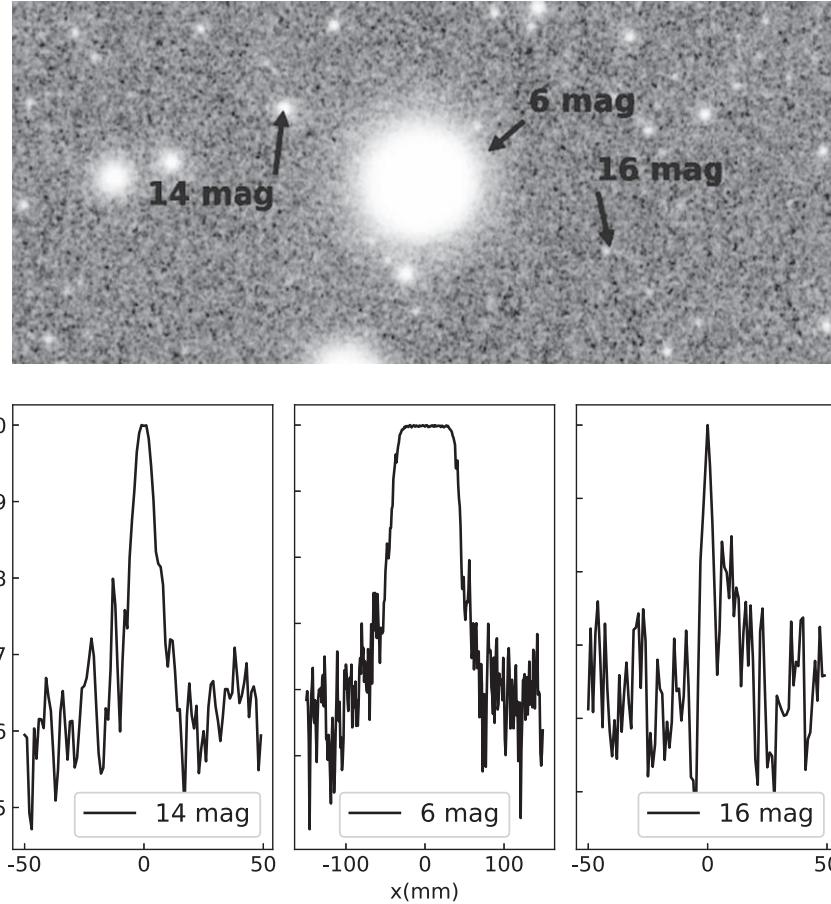


Figure 14. Intensity distributions of stars with different brightnesses, using plate BJ8304DA3822001 as an example. The upper panel shows the original digitized image of the plate. The lower panel, from left to right, displays the intensity distribution along a horizontal line passing through the center for 14, 6 and 16 magnitude objects, respectively. Each distribution has been normalized.

Acknowledgments

The authors express their sincere gratitude to NAOC, PMO, YNAO and QDO for their strong support and assistance in the preservation and transportation of plates. We also thank Professor Songzhu Lan, Professor Guohong Fu, Professor Baoan Yao, Professor Shuhe Wang, Professor Yaqing Mao, Professor Haibin Zhao, Professor Jiebing Yang and Professor Jianwei Zhang for their diligent work at organizing the plate information. The authors thank Professor Junliang Zhao, Professor Bochen Qian and Professor Jiaji Wang for their guidance on plate quality grading. The authors appreciate the advice provided by Professor Wenjing Jin, Professor Yongheng Zhao, Professor Rongchuan Qiao, Professor Zi Zhu and Professor Zhenyu Wu on plate image processing. The authors thank Youfen Wang, Yihan Tao, Linying Mi and Jun Han for their work on the database. Additionally, we express our sincere gratitude to Professor Hugh Jones for his detailed advice on the content of the article. This work is supported by the Shanghai Science and Technology Innovation Action Plan (grant No. 21511104100), the Global Common Challenge

Special Project (grant No. 018GJHZ2023110GC) and the China National Key Basic Research Program (grant No. 2012FY120500).

ORCID iDs

Shi-Yin Shen (沈世银) <https://orcid.org/0000-0002-3073-5871>

Quan-Feng Xu (徐权峰) <https://orcid.org/0000-0002-9572-6212>

References

Bertin, E., & Arnouts, S. 1996, *A&AS*, 117, 393
 Brown, A. G. A., Vallenari, A., Prusti, T., et al. 2018, *A&A*, 616, A1
 Bu, Y., Chen, F., & Pan, J. 2014, *NewA*, 28, 35
 Chauhan, V. K., Dahiya, K., & Sharma, A. 2019, *Artif. Intell. Rev.*, 52, 803
 Cleveland, W. 1979, *J. Am. Stat. Assoc.*, 74, 829
 De Cuyper, J., & Winter, L. 2005, in *ASP Conf. Ser.* 347, *Astronomical Data Analysis Software and Systems XIV*, 347, ed. P. Shopbell, M. Britton, & R. Ebert (San Francisco, CA: ASP), 651
 Fu, C., & Zhao, J. 2010, *Ann. Shanghai Obser. Acad. Sinica*, 00, 140
 Girard, T., Platais, I., Kozhurina-Platais, V., van Altena, W., & Lopez, C. 1998, *AJ*, 115, 855

Grindlay, J. E., & Griffin, R. E. 2012, in 285th Symp. Int.-Astronomical-Union, 285, ed. R. Elizabeth Griffin, Robert J. Hanisch, & Robert L. Seaman (Cambridge: Cambridge University Press), 243

Jin, W., Tang, Zhenghong and Wang, S., Yao, B., et al. 2007, *Prog. Astron.*, 25, 1

Lang, D., Hogg, D. W., Mierle, K., Blanton, M., & Roweis, S. 2010, *AJ*, 139, 1782

Ma, Z., Li, W., Wang, L., Xu, H., & Zhu, J. 2016, in IEEE 13th Int. Conf. on Signal Processing (ICSP), ed. Baozong Yuan, Quiqi Ruan, Yao Zhao, & Gaoyun An (New York: IEEE), 1021

Noble, W. S. 2006, *NatBi*, 24, 1565

Pedregosa, F., Varoquaux, G., Gramfort, A., et al. 2011, *J. Mach. Learn. Res.*, 12, 2825

Prusti, T., de Bruijne, J. H. J., Brown, A. G. A., et al. 2016, *A&A*, 595, A1

Robert, V., Desmars, J., Lainey, V., et al. 2021, *A&A*, 652, A3

Shupe, D., Moshir, M., Li, J., et al. 2005, in *ASP Conf. Ser.* 347, *Astronomical Data Analysis Software and Systems XIV*, ed. P. Shopbell, M. Britton, & R. Ebert (San Francisco, CA: ASP), 491

Simcoe, R. J., Grindlay, J. E., Los, E. J., et al. 2006, *Proc. SPIE*, 6312, 631217

Tuvikene, T., Edelmann, H., Groote, D., & Enke, H. 2014, in *Astroplate 2014*, ed. L. Mišková & L. Vítek (Prague: Institute of Chemical Technology), 127

Yang, M., Yu, Y., Wang, L., et al. 2024, *Astron. Tech. Instrum.*, 1, 1

Yu, Y., Zhao, J., Tang, Z., & Liao, S. 2013, *Astron. Res. Technol.*, 10, 288

Yu, Y., Zhao, J.-H., Tang, Z.-H., & Shang, Z.-J. 2017, *RAA*, 17, 28

Zheng, S., Liu, J., & Tian, J. 2005, *Pattern Recognit. Lett.*, 26, 147